Corrections to static antenna patterns in the kHz range

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Rube Goldberg

(A comically involved, complicated invention, laboriously contrived to perform a simple operation)

LIGO documents on the frequency dependence of the LIGO antenna patterns and the implication for calibration:

- T970101-B, D. Sigg, Strain calibration in LIGO,
- T030296, D. Sigg and R. Savage, Analysis proposal to search for gravitational waves at multiples of the LIGO arm cavity free-spectral-range frequency,
- T030186, J. Markowicz, R.L. Savage, and P. Schwinberg, *Development of a readout scheme for high-frequency gravitational waves*,
- G050205, M. Rakhmanov and R. Savage, LIGO detector response at high frequencies and its implications for calibration above 1kHz,
- T050136, Hunter Elliott, Analysis of the frequency dependence of the LIGO directional sensitivity (Antenna Pattern) and implications for detector calibration,
- G060665, M. Rakhmanov, Frequency corrections to antenna-patterns: forward detector transfer function,
- G060667, R. Savage et al., LIGO high-frequency response to length- and GW-induced optical path length variations,
- T070037, Jeffrey Parker, Development of a high-frequency burst analysis pipeline.

Antenna patterns at f = 0 or DC:

 $F_+(\phi, \theta)$ – response to +polarization,

 $F_{\times}(\phi, \theta)$ – response to ×polarization. [* Is there something like $F(\phi, \theta, \psi)$ *] ?



Long time ago ...



T970101-B, D.Sigg, Strain calibration in LIGO.

$H_{gw}(s)$: it's as simple as one, two, three...

Source location: (ϕ, θ) , polarization angle: ψ , unit vector \vec{n} , and rotation matrix R:

$$\begin{aligned} n_x &= \sin \theta \cos \phi \\ n_y &= \sin \theta \sin \phi , \\ n_z &= \cos \theta \end{aligned} \qquad E = \begin{pmatrix} 1 & 0 & 0 \\ 0 & -1 & 0 \\ 0 & 0 & 0 \end{pmatrix}, \qquad \begin{aligned} R &= R_z(\psi) R_y(\theta) R_z(\phi), \\ E' &= R^T E R \end{aligned}$$

Introduce the equivalent phase response and the Fabry-Perot cavity field response ($s = 2\pi i f$):

$$\phi_i = \frac{A_i - B_i e^{-2sT}}{2sT}, \qquad \qquad H_{\rm FP} = \frac{1 - r_a r_b}{1 - r_a r_b e^{-2sT}},$$

where r_a, r_b are mirror reflectivities and A_i and B_i (i = x, y) are phase factors:

$$A_i = \frac{1 - e^{-(1 - n_i)sT}}{1 - n_i}, \qquad B_i = \frac{1 - e^{-(1 + n_i)sT}}{1 + n_i}.$$

Then the response to gravitational waves is

$$H_{gw} = \frac{1}{2} H_{\rm FP} \left(E'_{xx} \phi_x - E'_{yy} \phi_y \right).$$

For a polarized source, we can also introduce the redundant responses:

$$H_{+} = H_{gw}|_{\psi=0}, \qquad H_{\times} = H_{gw}|_{\psi=45^{\circ}}, \qquad H_{circ} = H_{+} + iH_{\times}.$$



Magnitude of $H_{gw}(s)$

Collect all angles: $\Omega = (\phi, \theta, \psi).$

Detector response = convolution:

$$x(t) = \int_0^T \left[H_+(t - t', \Omega) \ h_+(t') + H_\times(t - t', \Omega) \ h_\times(t') \right] \ dt'.$$

In Fourier domain:

$$\tilde{x}(f) = H_+(f,\Omega) \ \tilde{h}_+(f) + H_\times(f,\Omega) \ \tilde{h}_\times(f)$$

At low frequencies ($f \ll FSR$):

$$H_{+,\times}(f,\Omega) \approx F_{+,\times}(\Omega) * H_{\text{pole}}(f),$$

where $F_{+,\times}(\Omega)$ are static antenna-patterns.

 $H_{\rm pole}(f)$ is an approximate frequency response for optimal orientation:

$$H_{\text{pole}}(f) = \frac{1}{1 + if/f_{cav}}, \qquad f_{cav} \approx 86 \text{ Hz}.$$

This is called the long-wavelength approximation.

 $H_{gw}(s)$: exact and approximate forms. (source coordinates: $\phi = 0, \ \theta = 0, \ \psi = 0$).



 $H_{gw}(s)$: exact and approximate forms. (source coordinates: $\phi = 0, \ \theta = 20^{\circ}, \ \psi = 0$).



Comparison of $H_{gw}(s)$ and $H_L(s)$.

Calibration measures $H_L(f)$ and not $H_{gw}(f)$.



At low frequencies the magnitude of the length response and that of the gravitational-wave response are almost the same. The phase is slightly different though.



Complete response contains three effects.

	geometric factor	photon transit time	Fabry-Perot effect
approx.	$F_{+,\times}(\phi,\theta)$	1	$H_{\text{pole}} = \frac{1}{1 - s/s_{cav}}$
exact	$F_{+,\times}(\phi,\theta)$	sinc-function	$H_{\rm FP} = \frac{1 - r_a r_b}{1 - r_a r_b e^{-2sT}}$
Grishchuk	$F_{+,\times}(\phi,\theta)$	sinc-function	1

from D. Baskaran and L.P. Grishchuk, Classical and Quantum Gravity, Vol. 21, p. 4041, 2004

Essential steps:

- 1. begin with the exact formula: $H_{gw}(s)$,
- 2. take out the Fabry-Perot effect. (i.e. put $H_{\rm FP}(s) = 1$),
- 3. expand the result in powers of f (frequency) or s,
- 4. keep the zeroth and the first order terms,

The zeroth order terms represent static antenna patterns $[F_{+,\times}(\Omega)]$ and are called "electric components". The first order terms represent a correction from the frequency dependence of the antenna patterns (sinc-functions) and are called "magnetic components".

Grishchuk at CaJAGWR: "...the magnetic component provides a correction of up to 10% in the frequency band of 1200 Hz."

Take a close look.



Give me ten!



Can you give me more?





Put the cavity back, please!



Etimology of *Gravi-magnetic* component: It was suggested by an analogy with electro-magnetism, or the Lorentz force, because of the time derivative. But the actual coupling is $x \times dh/dt$ and not $dx/dt \times h$. From a more fundamental point of view: can gravity couple to velocity?

So "What's in a name?"

Lee Iacocca: "Everything!"

Conclusions:

- Antenna patterns are frequency dependent.
- The first order correction is good, but the exact formula is even better.
- The frequency dependence does not introduce significant corrections at frequencies in our bandwidth. (no need to change the calibration model.)
- Static antenna patterns combined with the single-pole transfer function well approximate the exact response. (two errors cancell each other!)